

Statistics, Source Detection, and Noise in the Optical / Infrared

ASTR 288C: Lecture 4

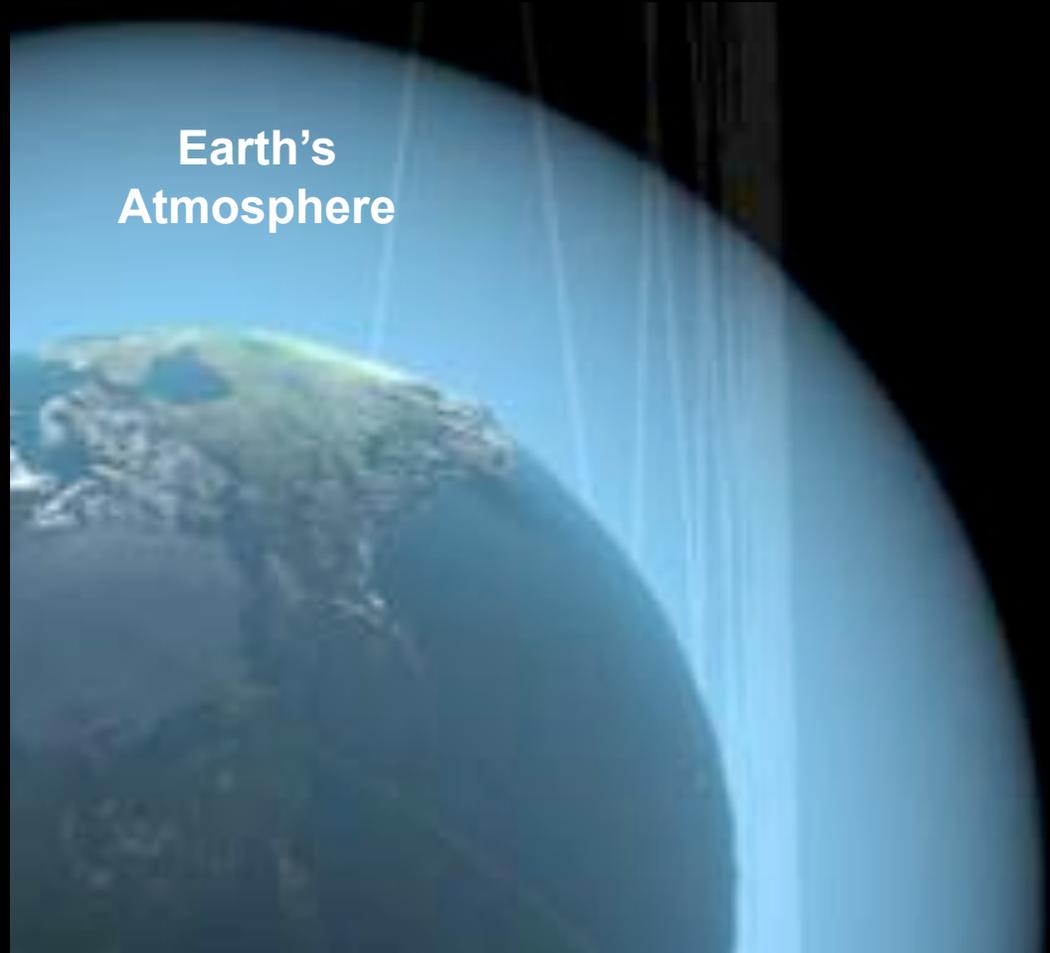
Sources of Detected Signal

Astrophysical
Sources



Atmosphere transmits most light at
optical wavelengths and within
“windows” in the infrared

Earth's
Atmosphere



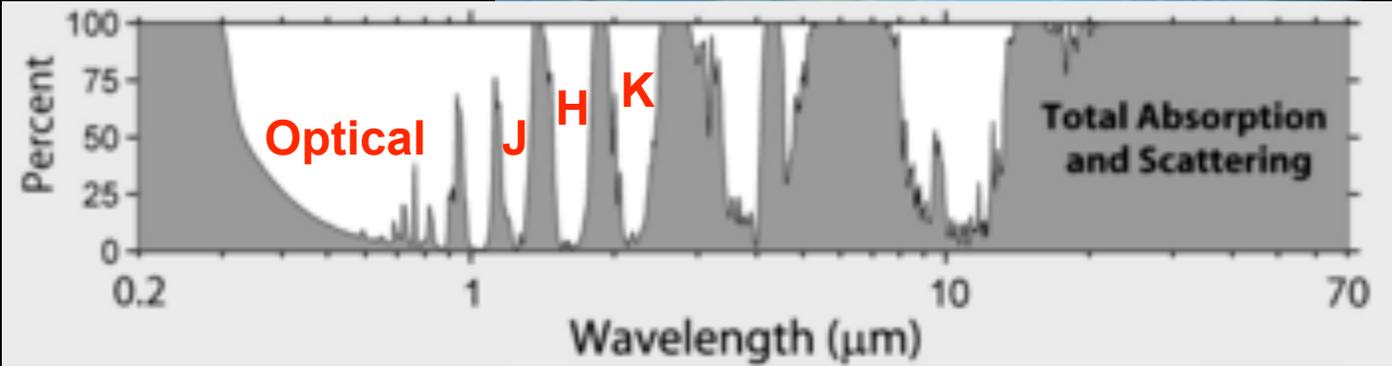
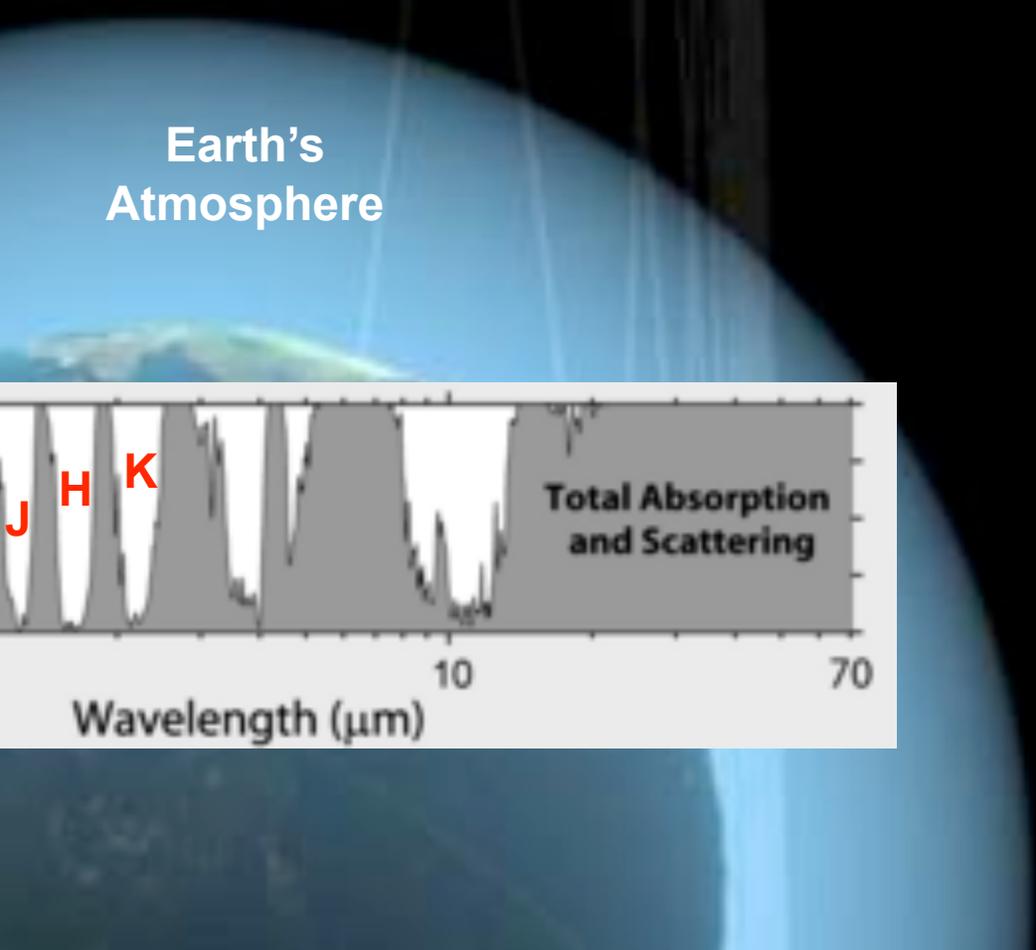
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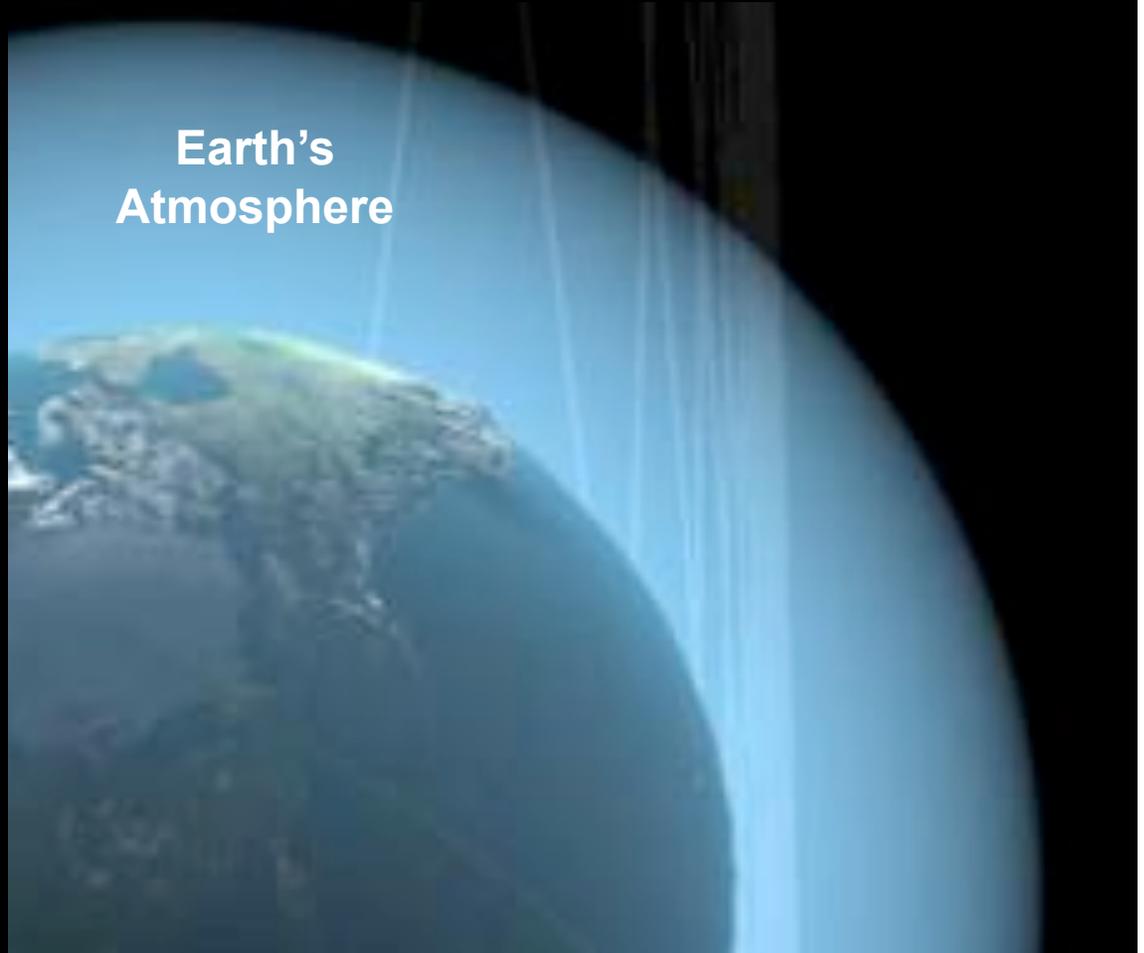
**Astrophysical
Sources**



**Atmosphere transmits most light at
optical wavelengths and within
“windows” in the infrared**

**BUT... it emits its own light,
especially in the infrared where this
emission is significantly greater
than from astrophysical sources**

**Earth's
Atmosphere**



Sources of Detected Signal

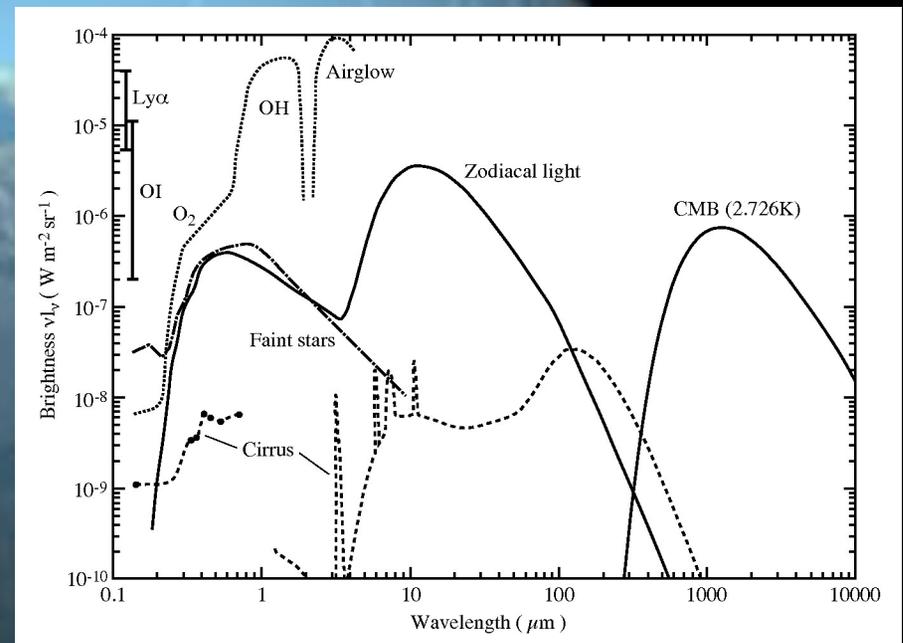
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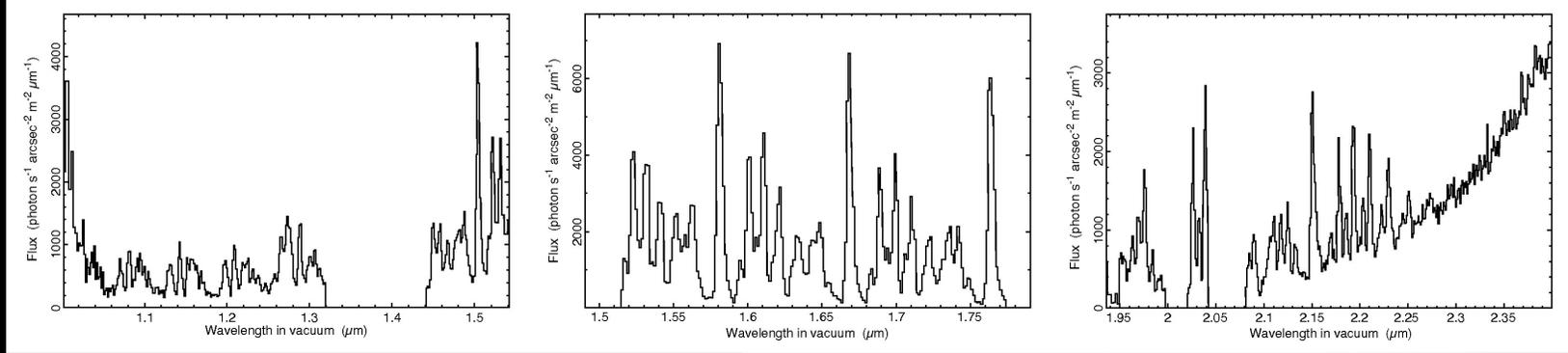
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Earth's Atmosphere

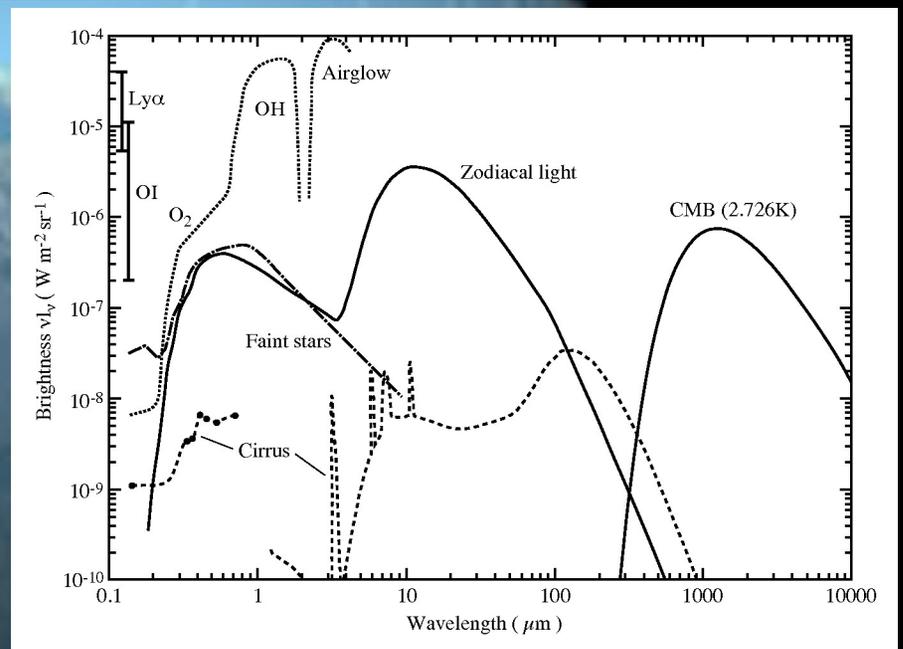


Sources of Detected Signal

Astrophysical Sources



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Astrophysical
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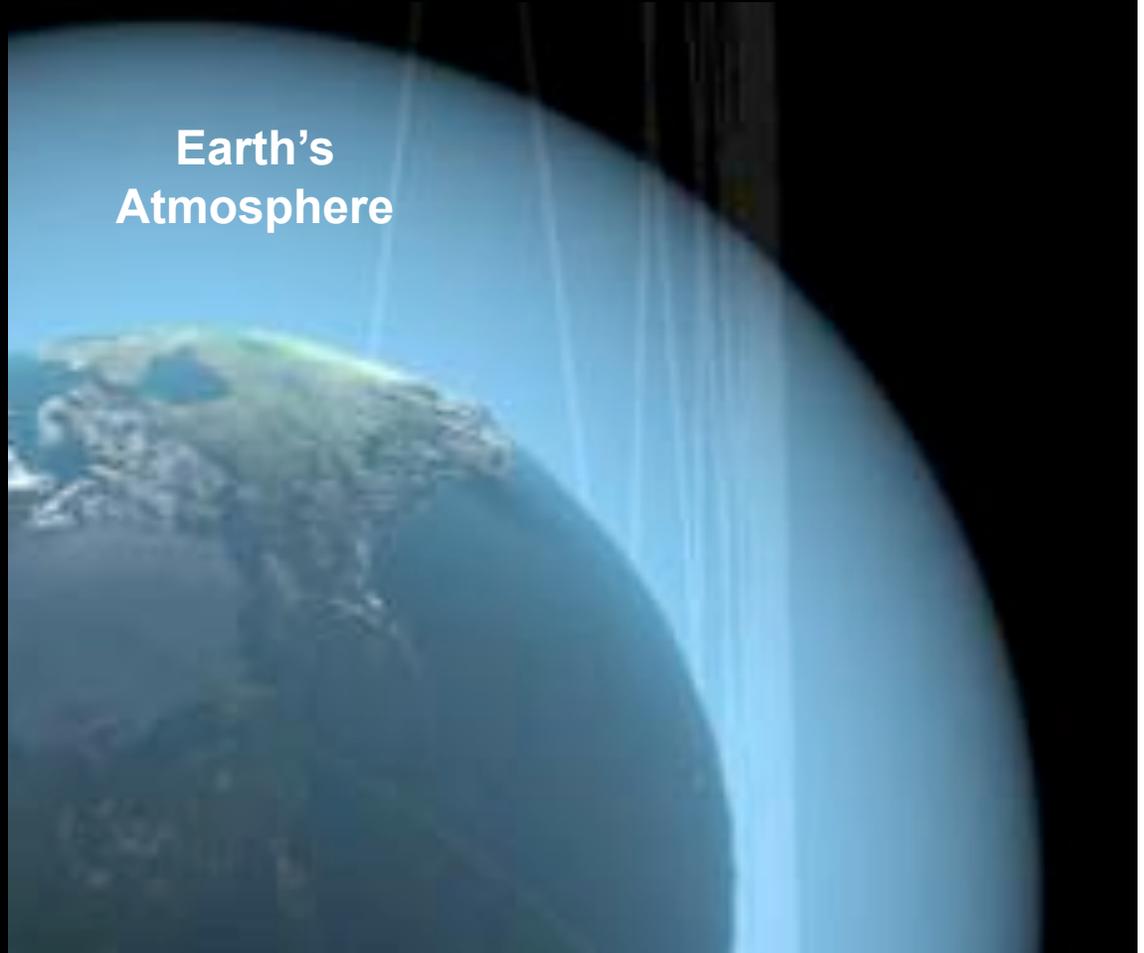


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BACKGROUND (“SKY”)

Earth’s
Atmosphere



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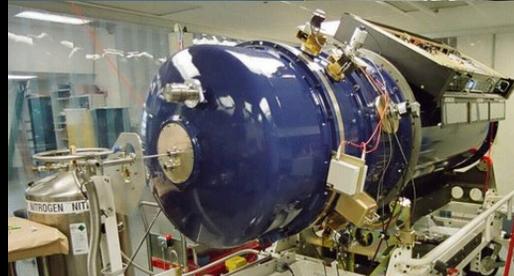
BACKGROUND (“SKY”)

Telescope and detector may be source of light reflections

Earth’s
Atmosphere

Telescope

Detector



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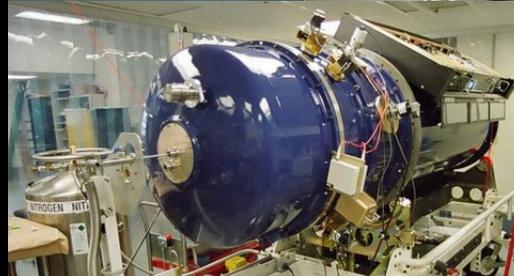
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Thermal emission of detector itself

Earth’s
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Telescope

Detector



Sources of Detected Signal

Astrophysical
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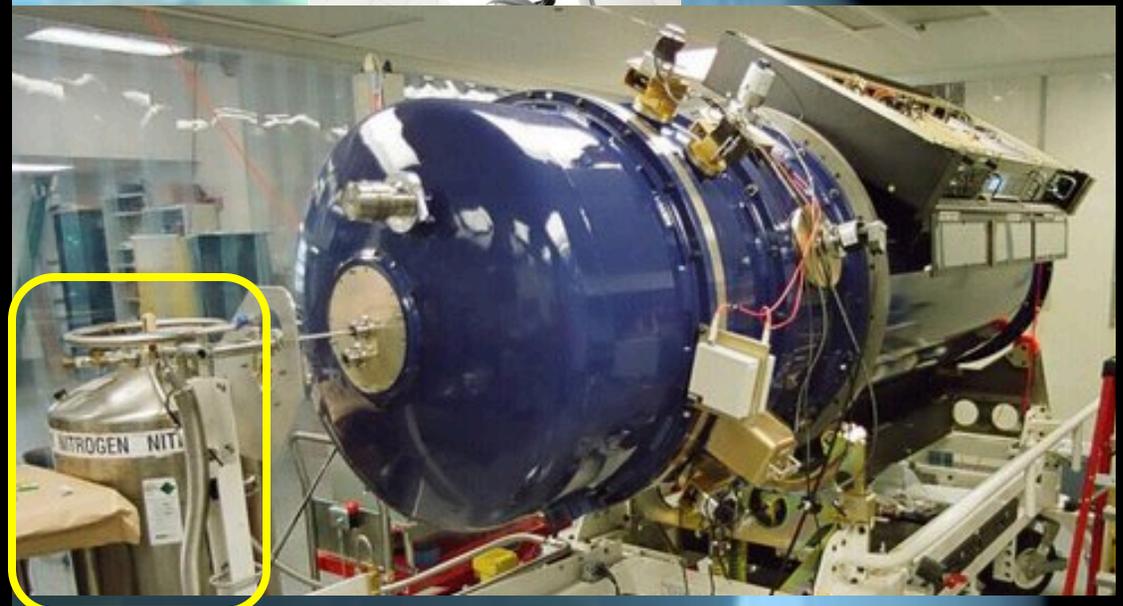
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(cool the detector)

Earth's
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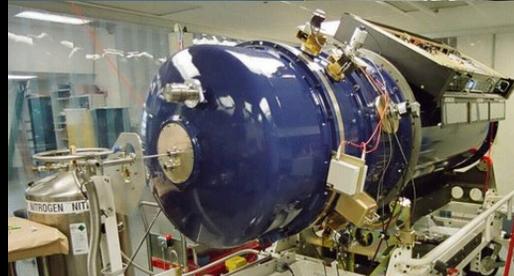
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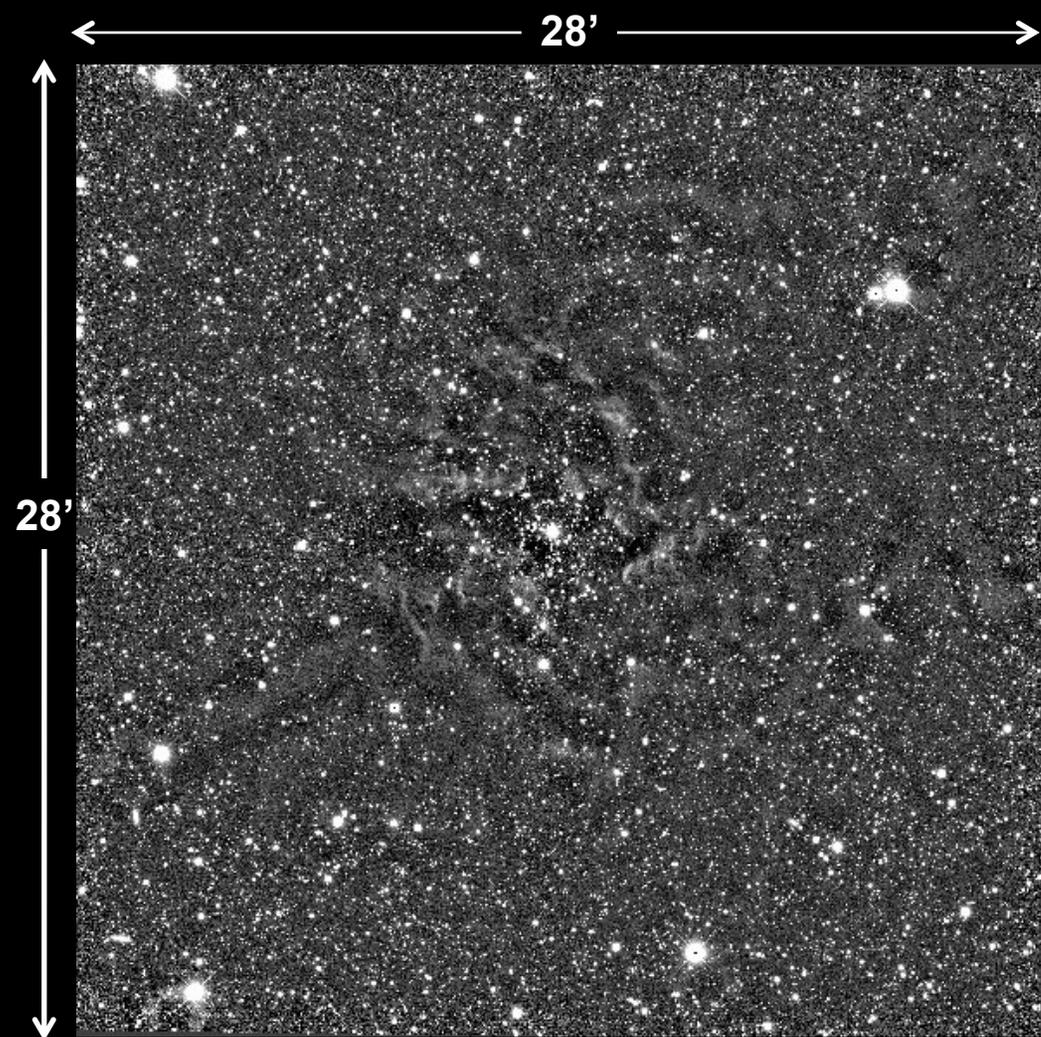
Earth's Atmosphere

Telescope

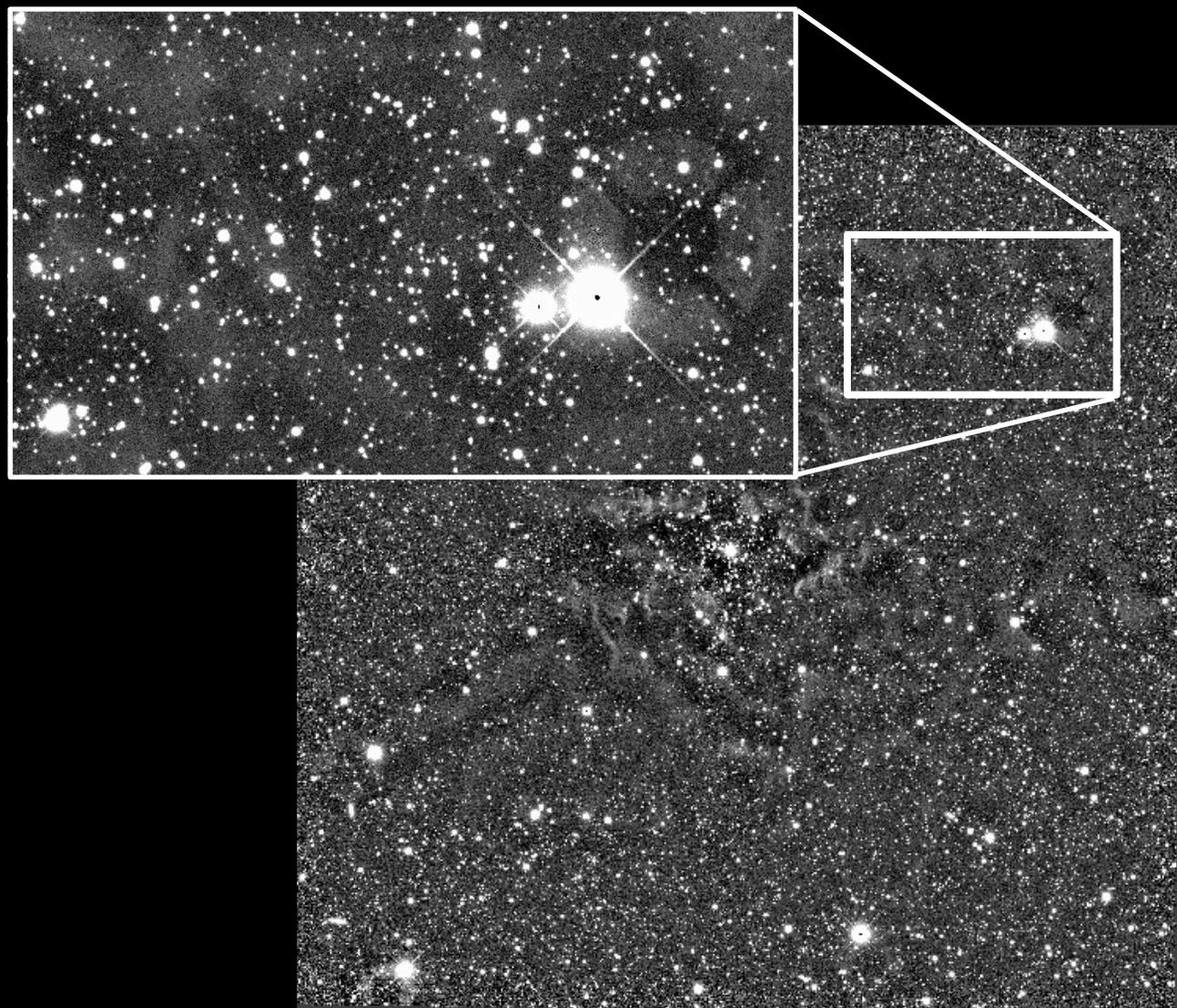
Detector



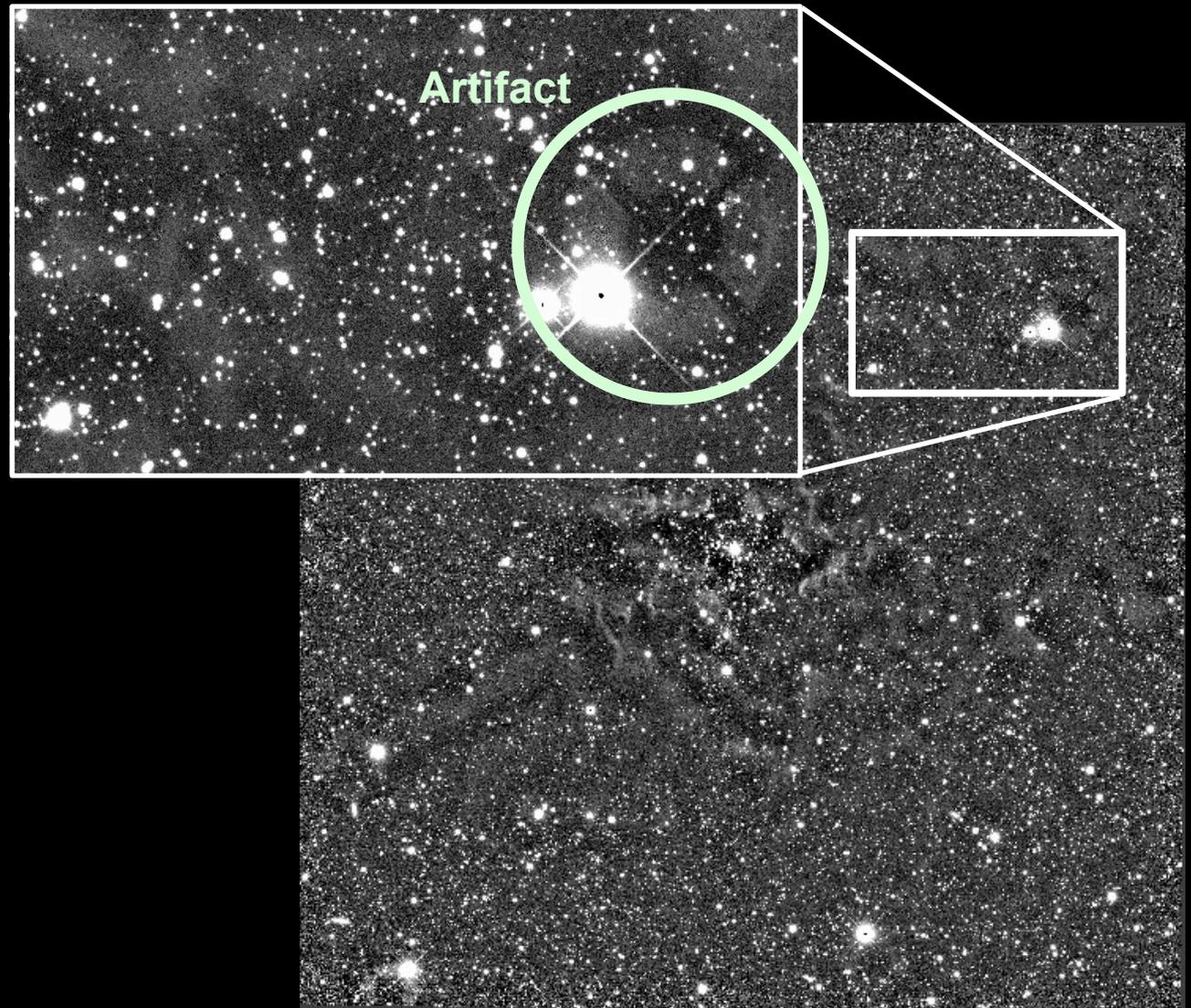
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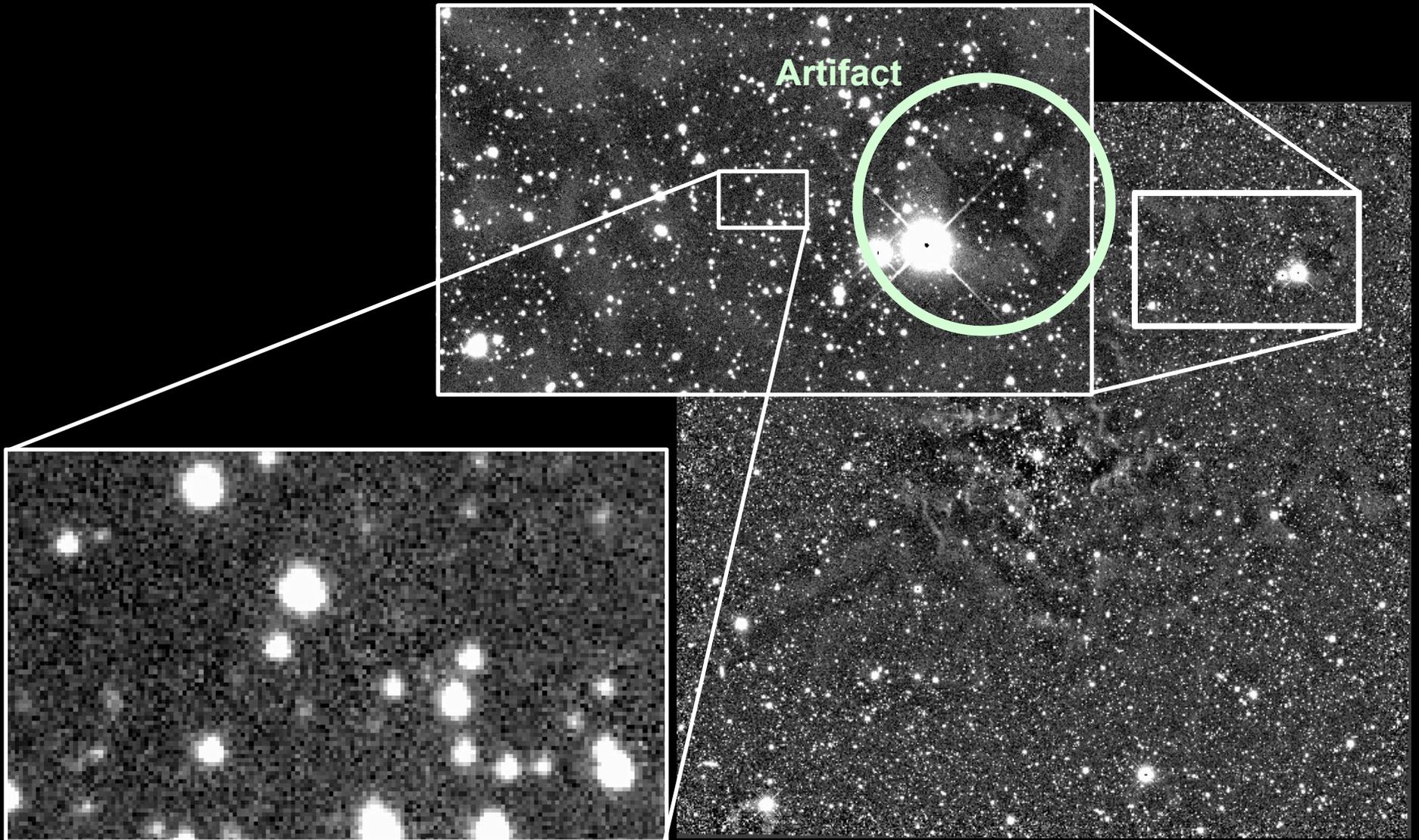
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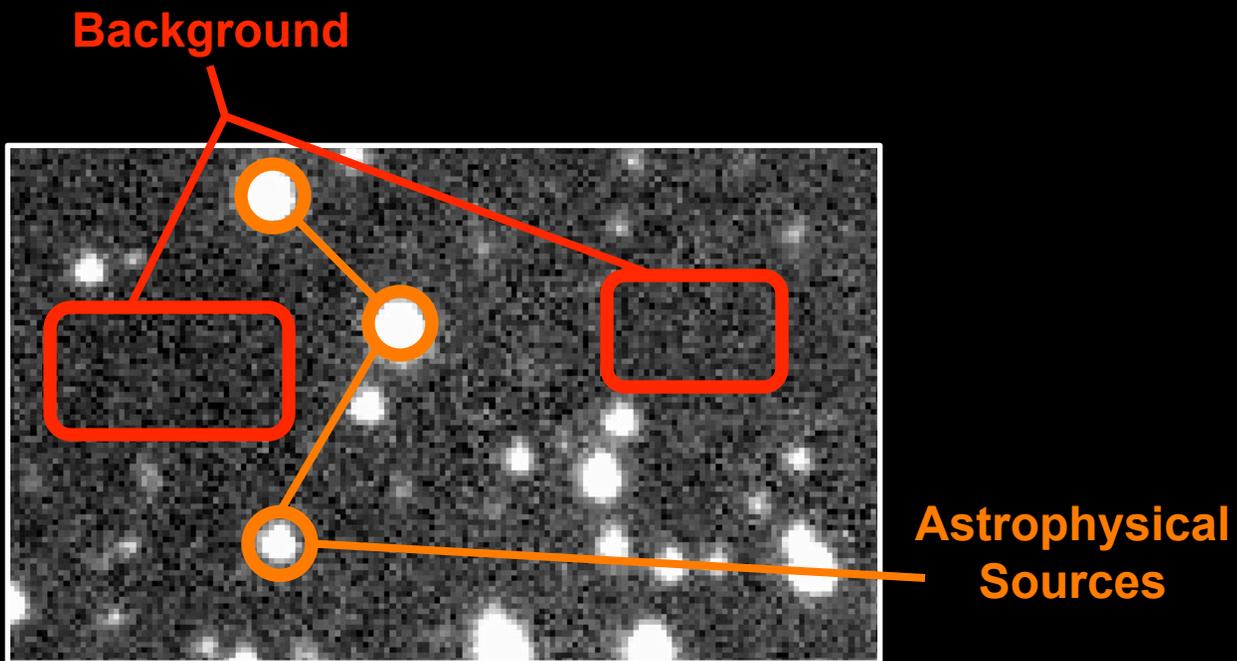
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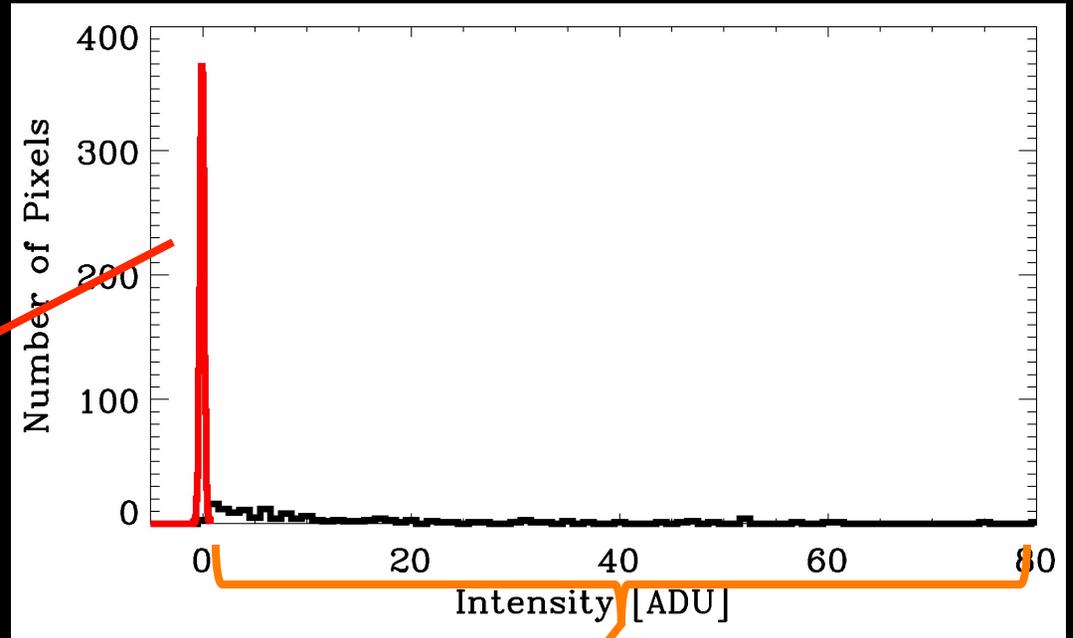
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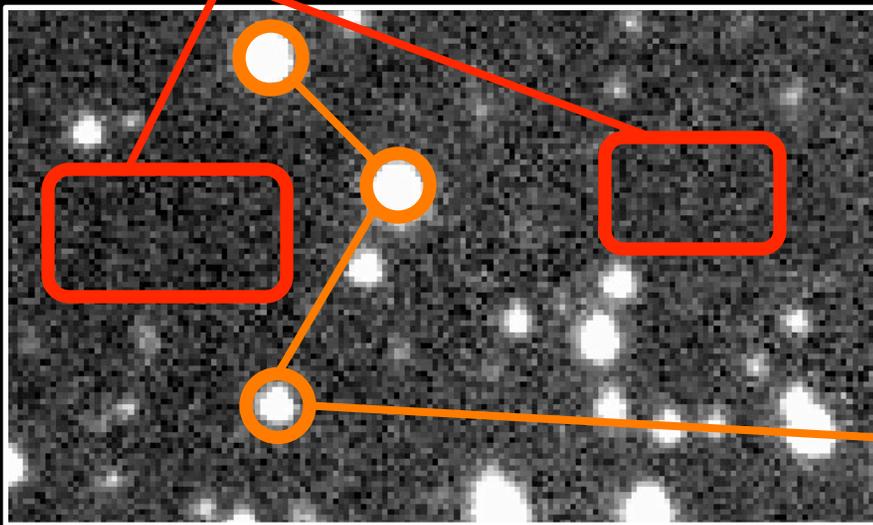
Sources of Detected Signal



Sources of Detected Signal



Background



Astrophysical Sources

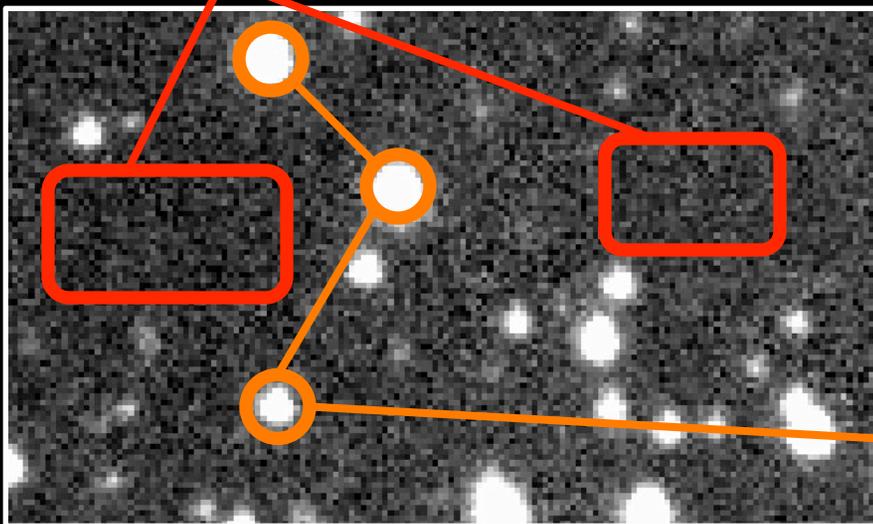
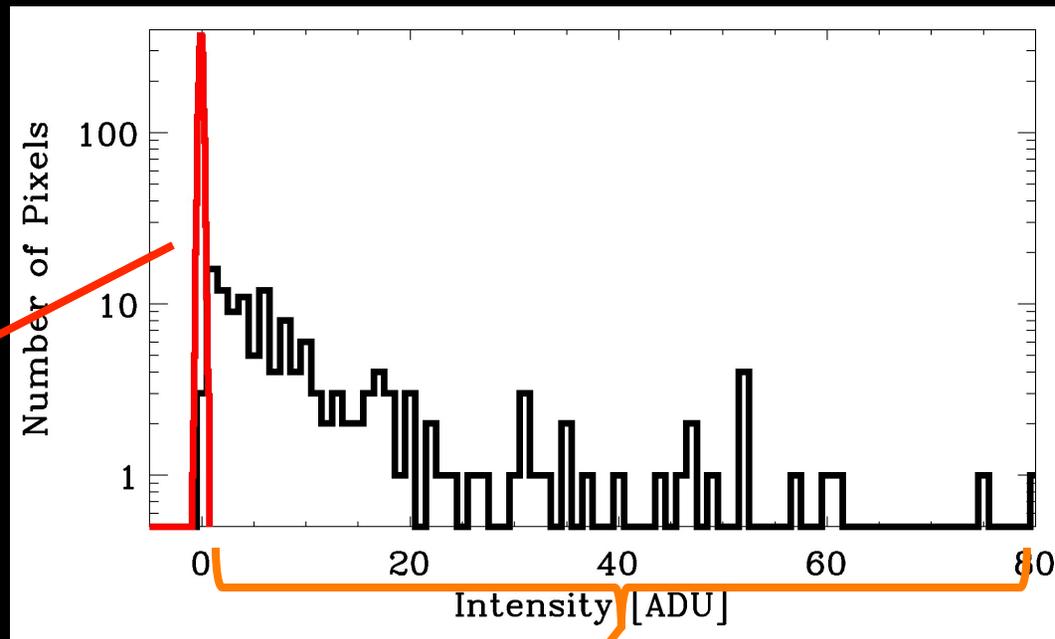
Sources of Detected Signal

Background ~ 0 ADU since these are REDUCED (not RAW) images.

REDUCED images have had the background level subtracted.

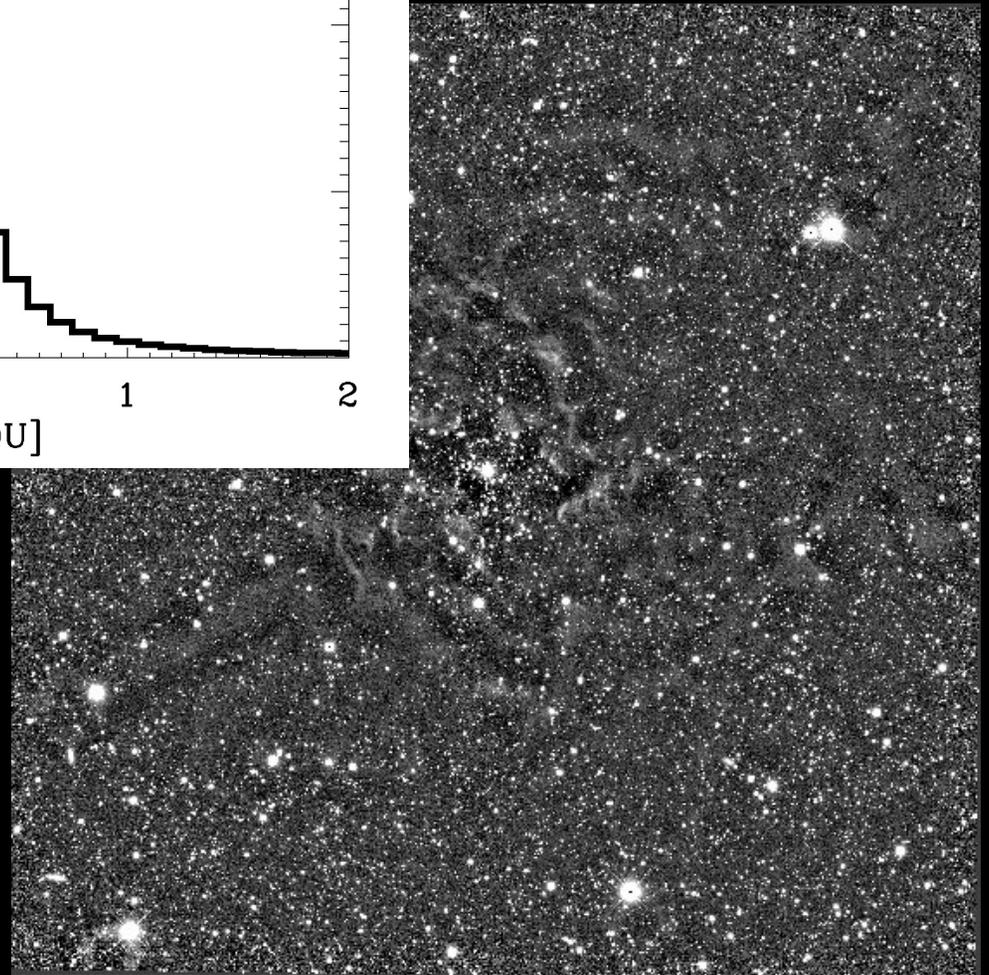
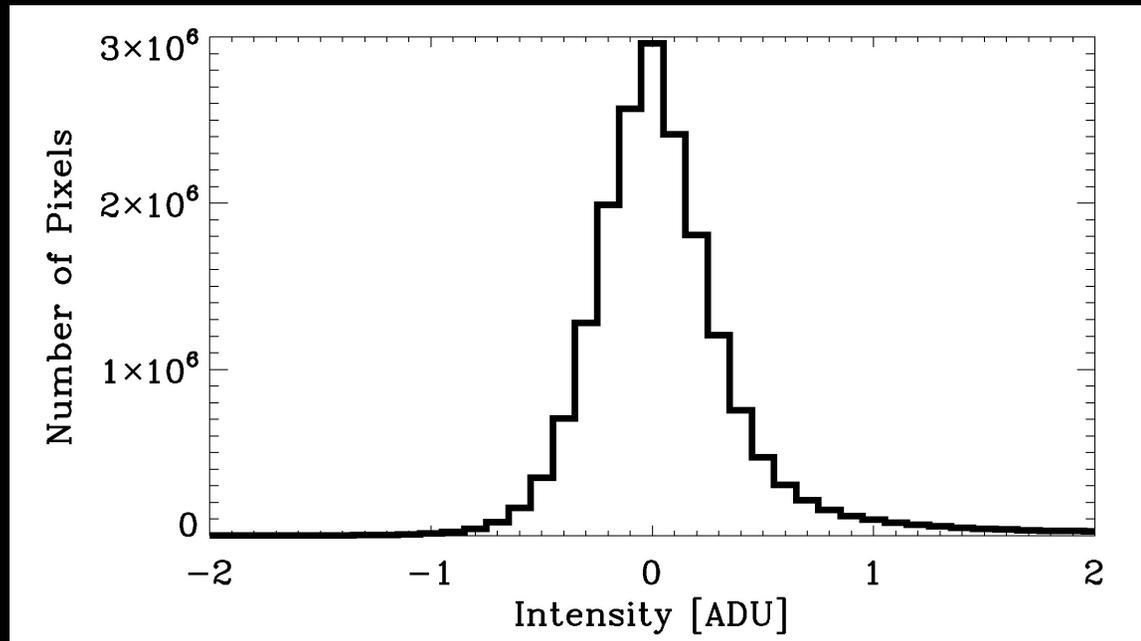
The RAW images have backgrounds of ~ 1000 ADU, significantly greater than the astrophysical sources identified below.

Background

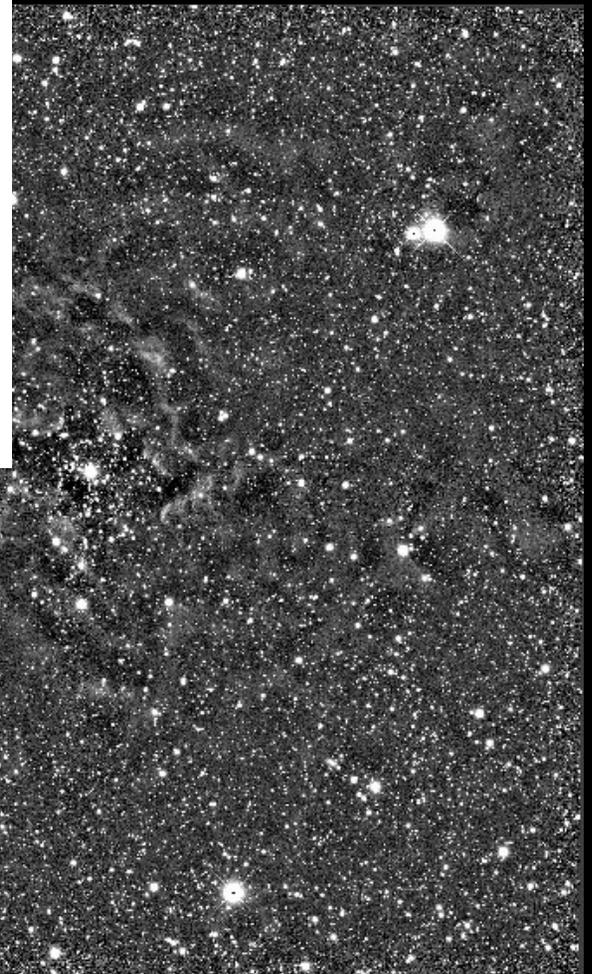
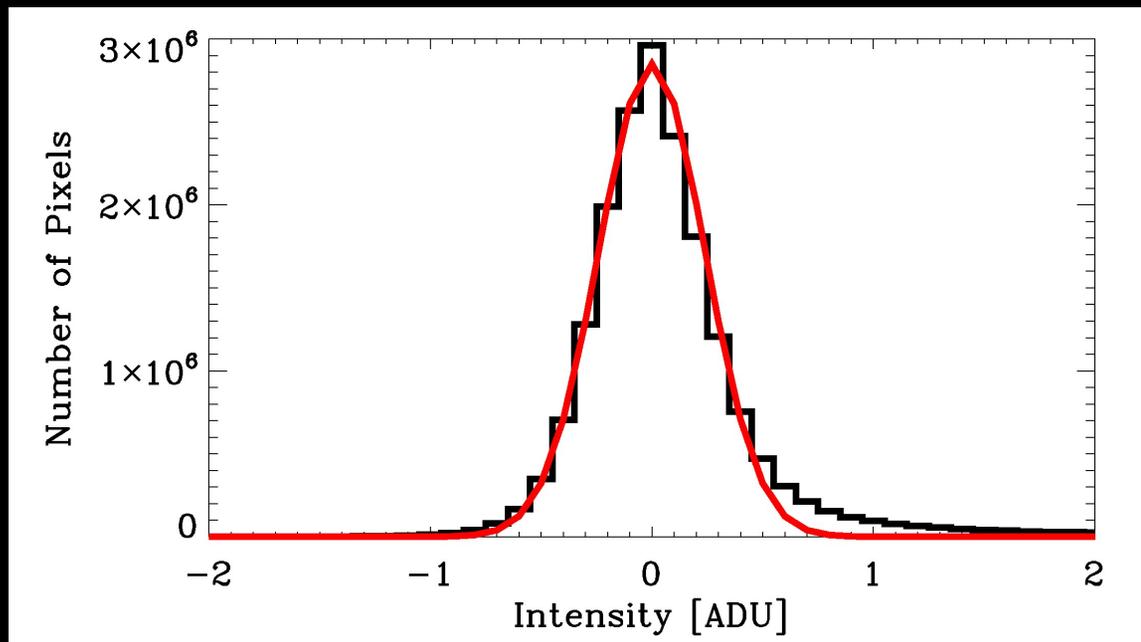


Astrophysical Sources

Sources of Detected Signal



Sources of Detected Signal

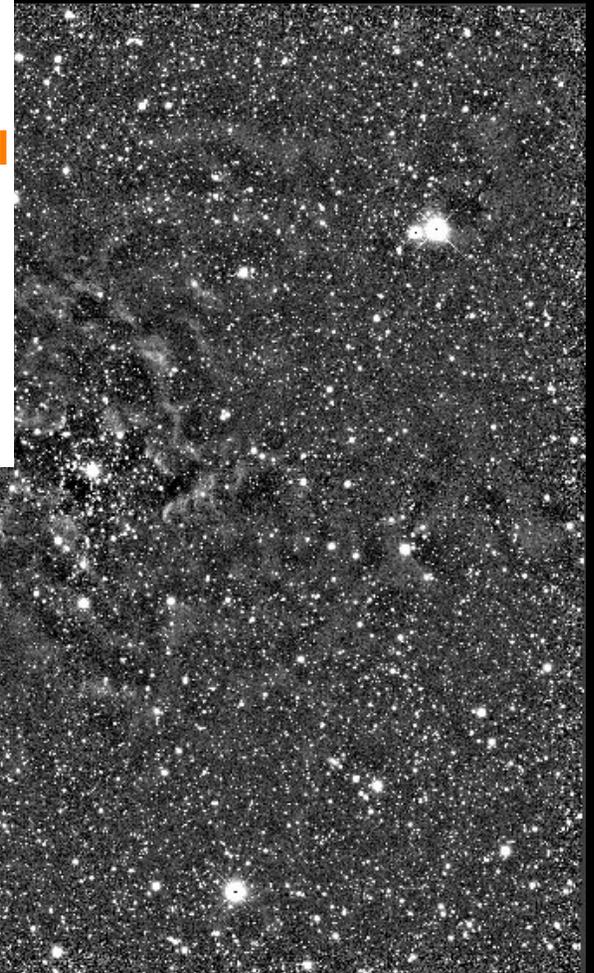
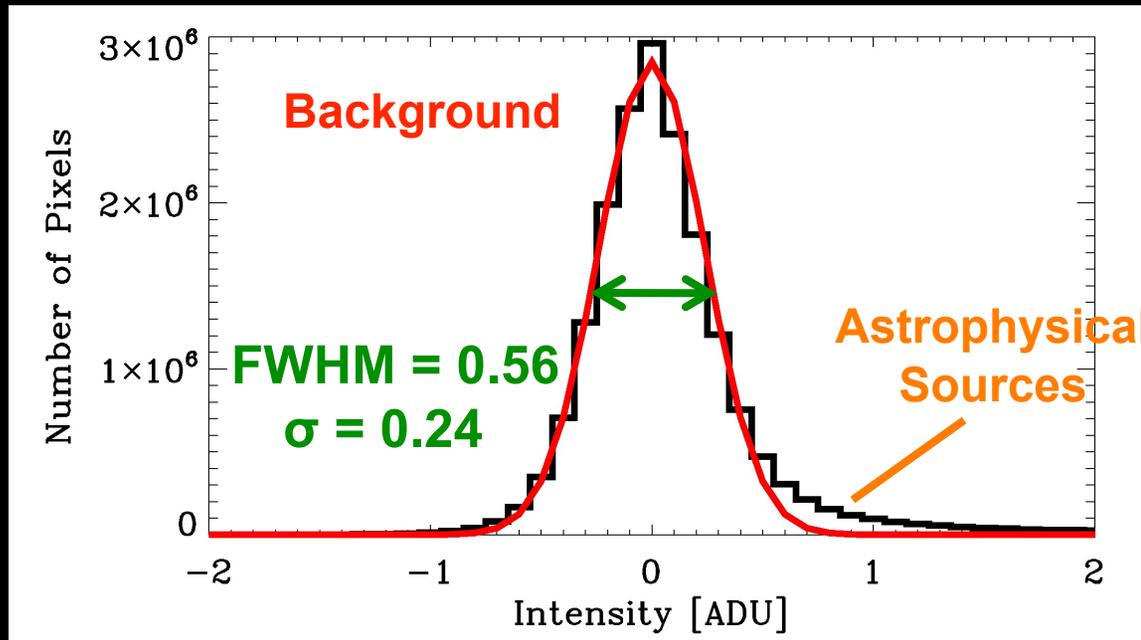


Gaussian

$$N = N_{\text{peak}} e^{-\frac{1}{2} \left[\frac{x - x_0}{\sigma} \right]^2}$$

$$\sigma = \frac{\text{FWHM}}{2 \sqrt{2 \ln 2}}$$

Sources of Detected Signal

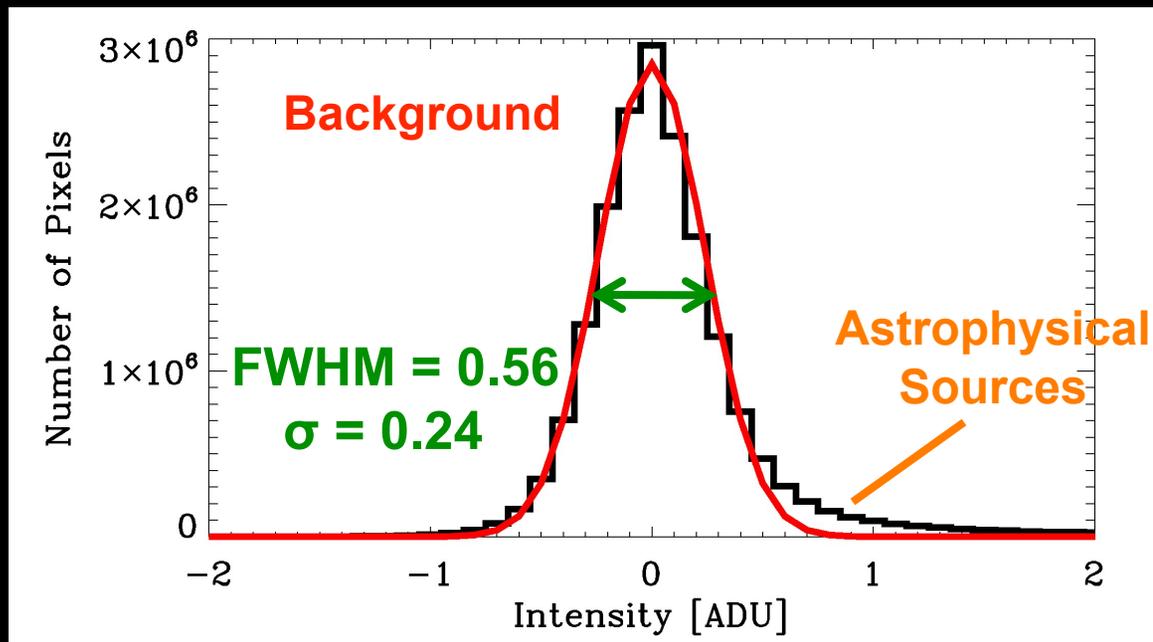


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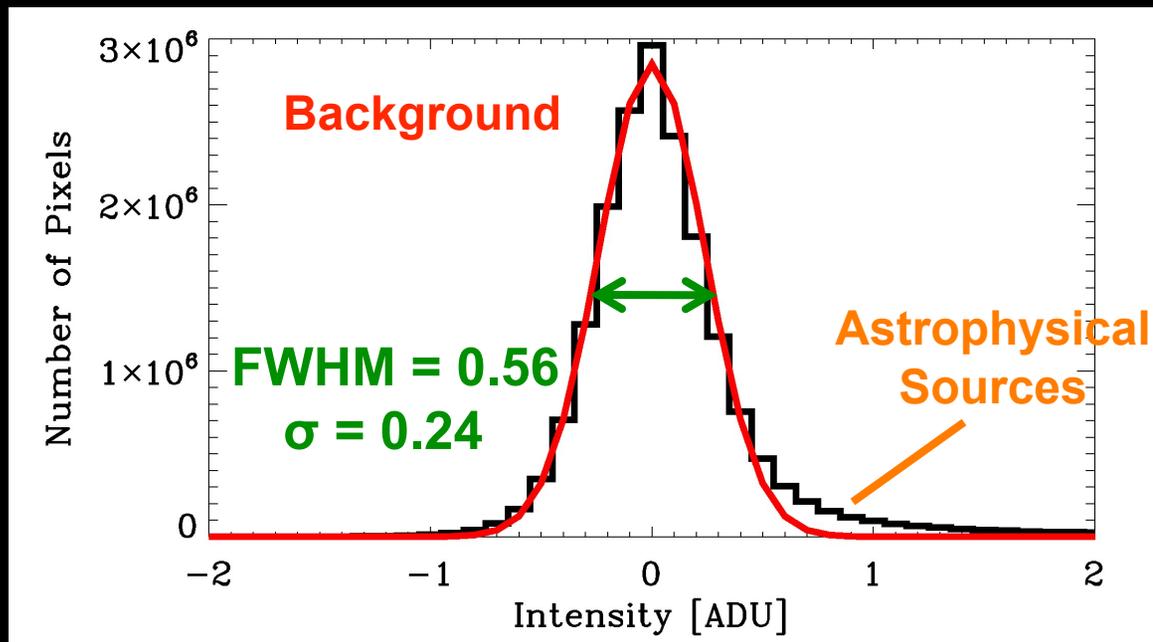
What is the “TYPICAL” background level?

MEAN intensity = 0.88 ADU

MEDIAN intensity = 0.02 ADU

MODE intensity = 0.00 ADU

Sources of Detected Signal



Gaussian

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What is the "TYPICAL" background level?

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What is meant by “TYPICAL”?

What is the typical weight of the pets in the neighborhood?

What is meant by “TYPICAL”?

What is the typical weight of the pets in the neighborhood?

MEAN = 9.9 lbs

MEDIAN = 10.0 lbs

MODE = 10.0 lbs



Buz
11 lbs



Dizzy
9 lbs



Tiger
10 lbs



Lanky
8 lbs



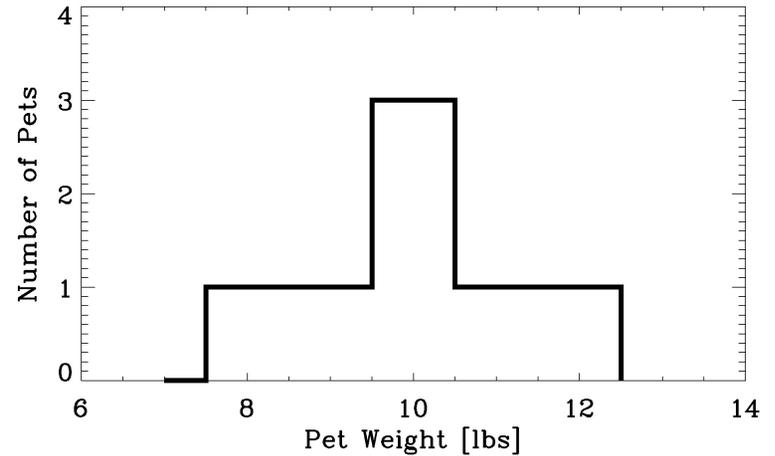
Dash
10 lbs



Fido
12 lbs



Teddy
10 lbs



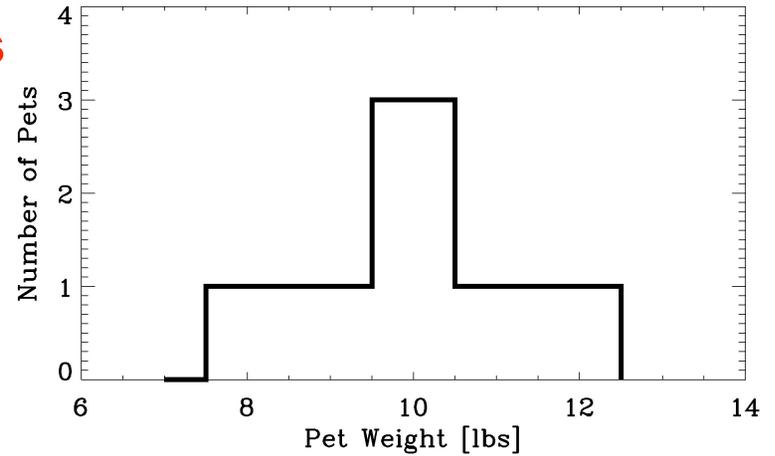
What is meant by “TYPICAL”?

What is the typical weight of the pets in the neighborhood?

MEAN = ~~970 lbs~~ 1008.6 lbs

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Buz
11 lbs



Dash
10 lbs



Dizzy
9 lbs



Fido
12 lbs



Elsa
8000 lbs



Tiger
10 lbs



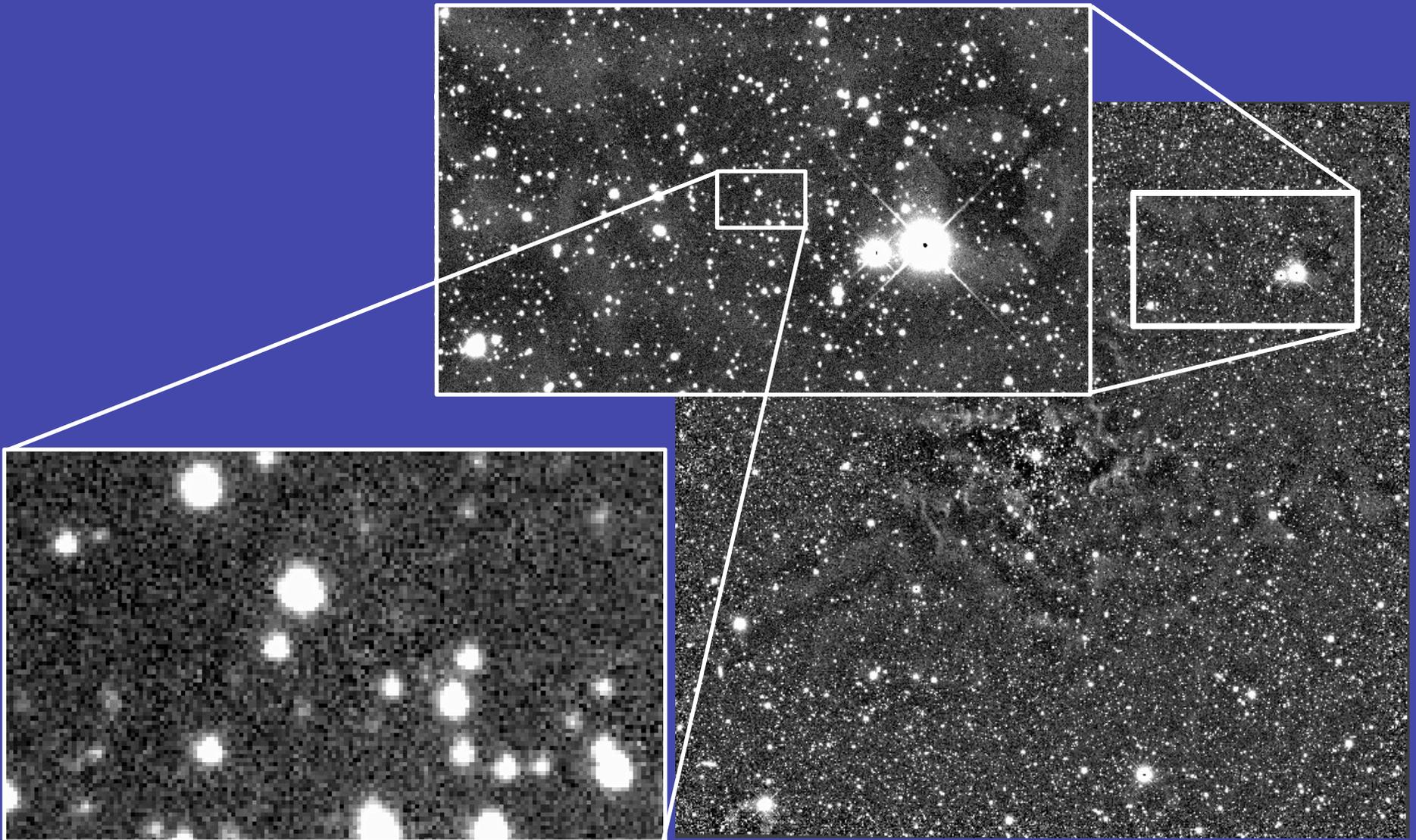
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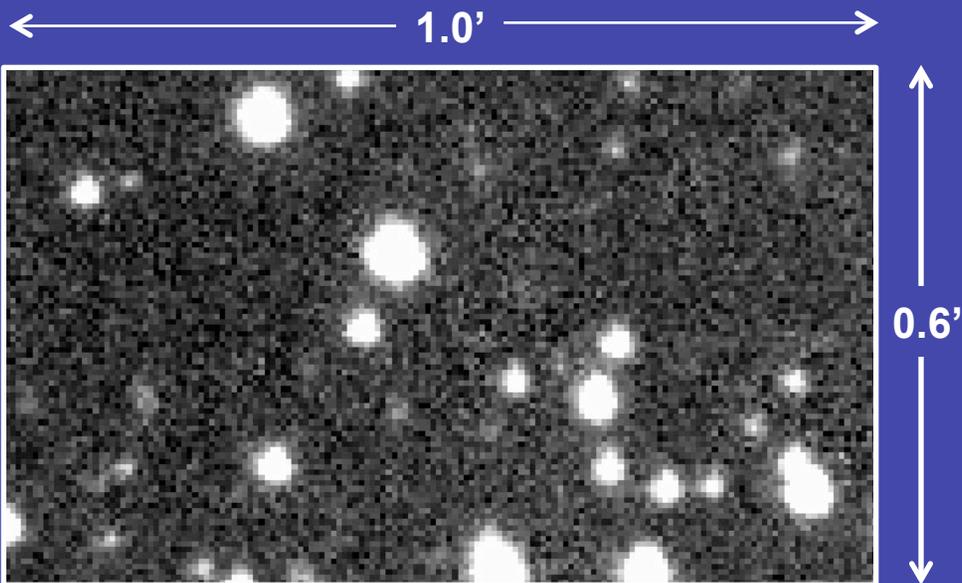
Source Detection

Identification of a real signal from an astrophysical source



Source Detection

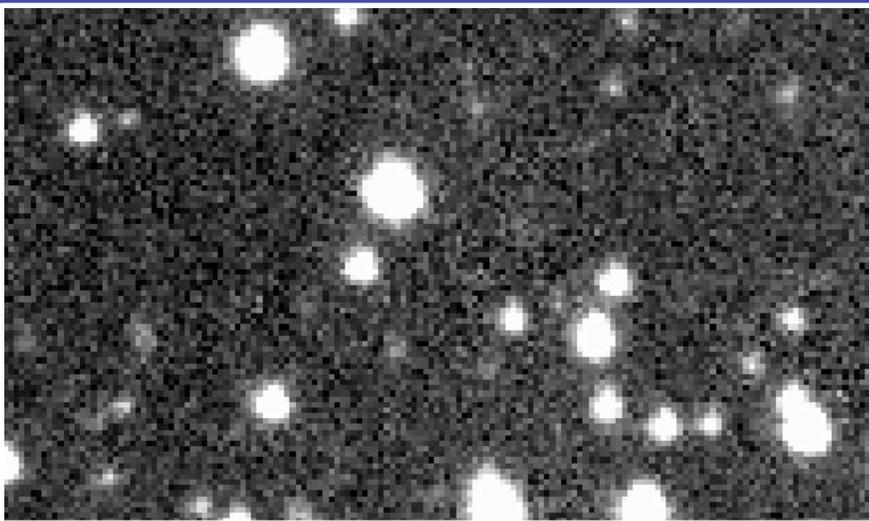
Identification of a real signal from an astrophysical source



Source Detection

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By eye, 31-36 sources in 1.0' x 0.6'



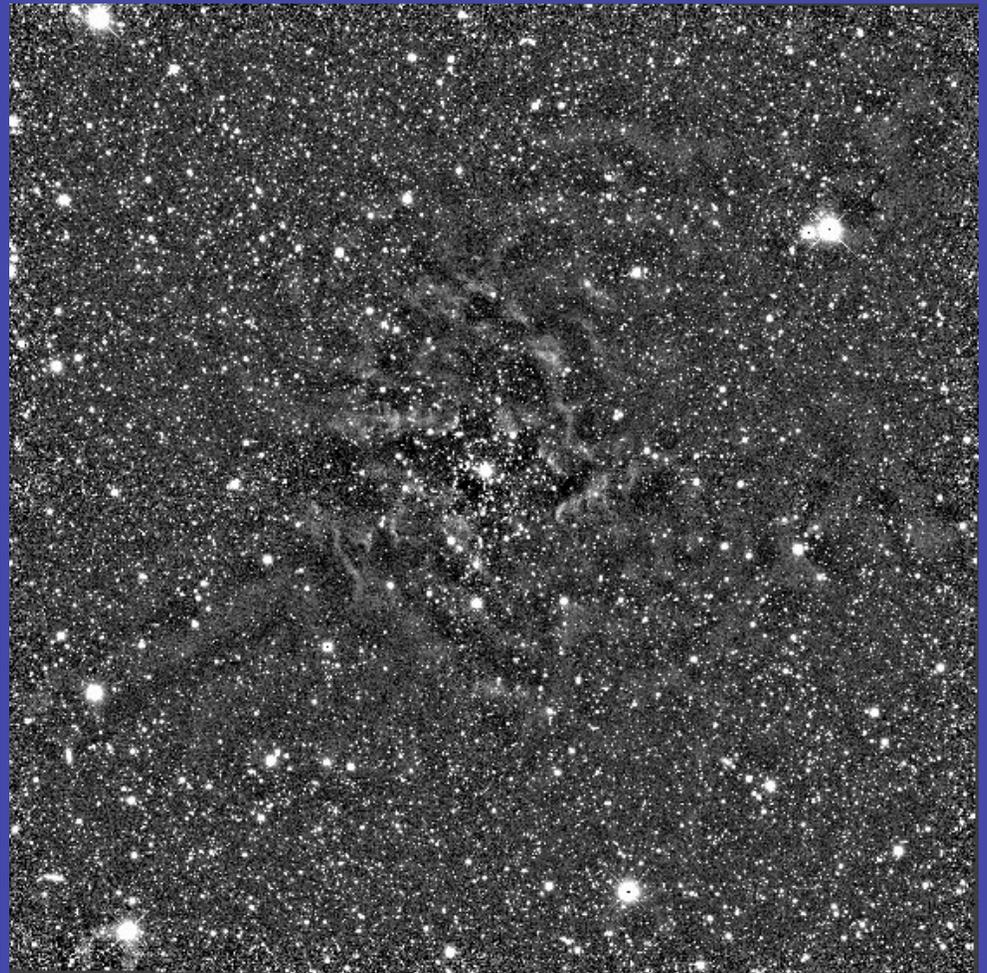
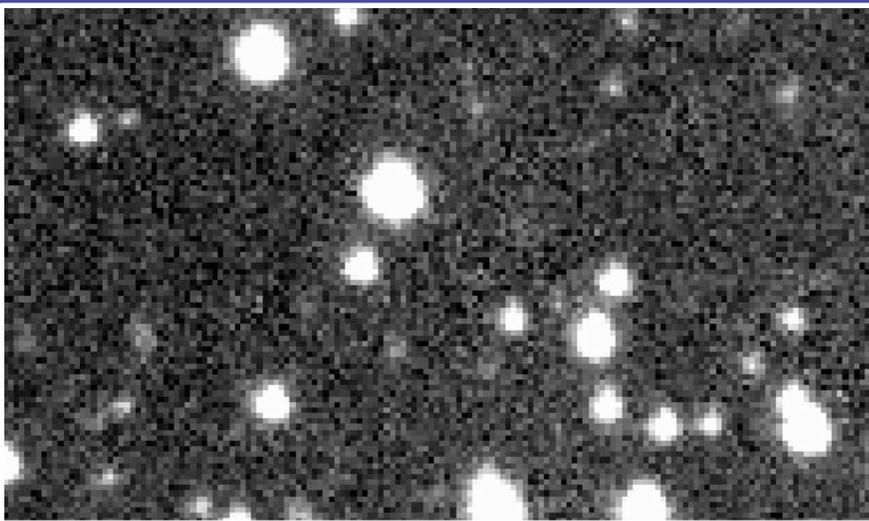
Source Detection

Identification of a real signal from an astrophysical source

ANY VOLUNTEERS ?!

Suggests ~43-50k sources in 28.8' x 28.8' !

By eye, 31-36 sources in 1.0' x 0.6'



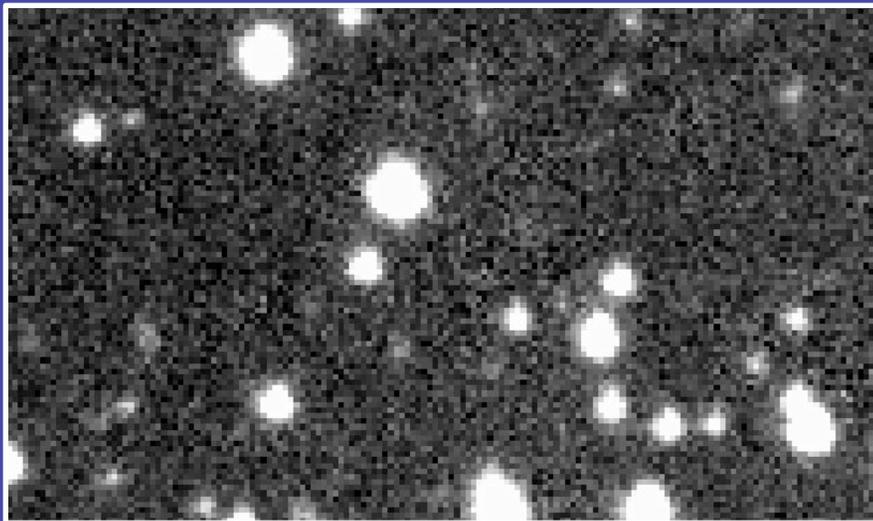
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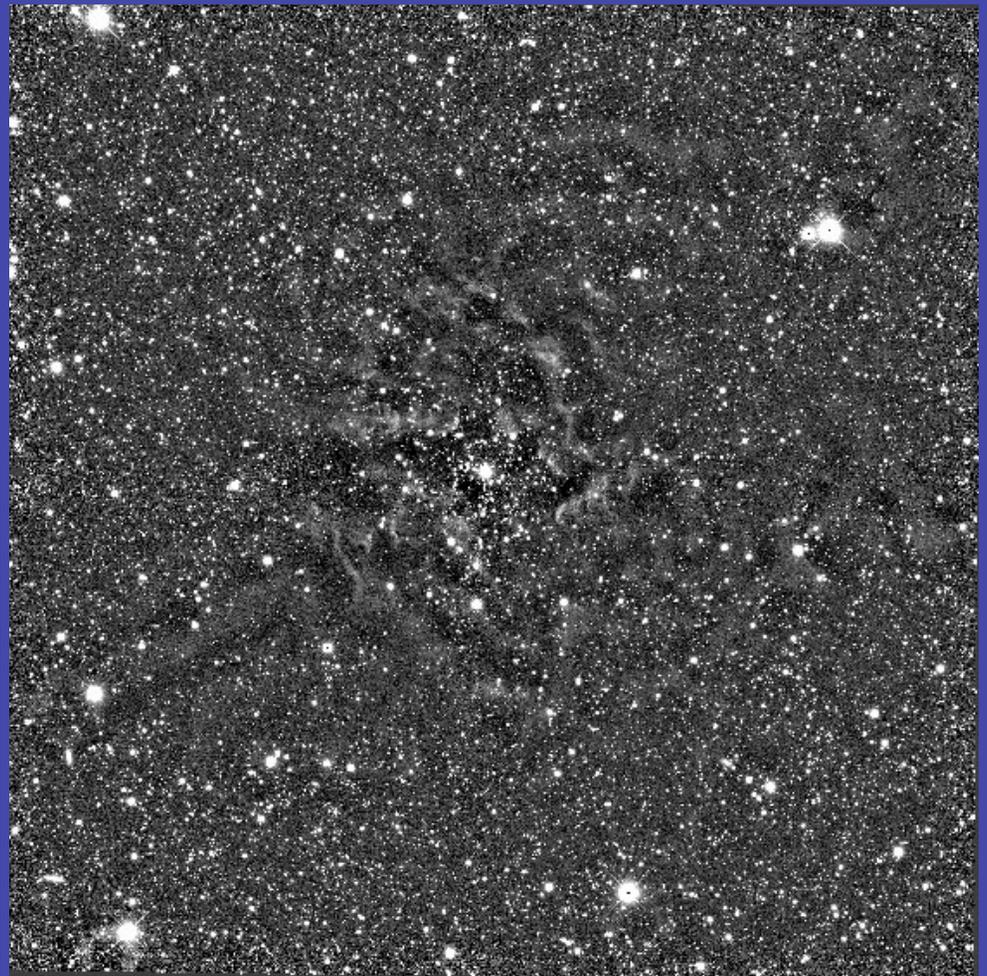
ANY VOLUNTEERS ?!

Clearly, we need an automated source detection method.

By eye, 31-36 sources in 1.0' x 0.6'



Suggests ~43-50k sources in 28.8' x 28.8' !



Source Detection

Identification of a real signal from an astrophysical source

Automated source detection procedures:

IRAF *daofind*

IDL *daofind*

SExtractor

